



# Metis Observations of Alfvénic Outflows Driven by Interchange Reconnection in a Pseudostreamer

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## Abstract

This study presents observations of a large pseudostreamer solar eruption and, in particular, the post-eruption relaxation phase, as captured by Metis, on board the Solar Orbiter, on 2022 October 12, during its perihelion passage. Utilizing total-brightness data, we observe the outward propagation of helical features up to 3 solar radii along a radial column that appears to correspond to the stalk of the pseudostreamer. The helical structures persisted for more than 3 hr following a jet-like coronal mass ejection associated with a polar crown prominence eruption. A notable trend is revealed: the inclination of these features decreases as their polar angle and height increase. Additionally, we measured their helix pitch. Despite the 2 minute time cadence limiting direct correspondence among filamentary structures in consecutive frames, we find that the Metis helical structure may be interpreted as a consequence of twist (nonlinear torsional Alfvén waves) and plasma liberated by interchange reconnection. A comparison was performed between the helix parameters as outlined by fine-scale outflow features and those obtained from synthetic white-light images derived from the high-resolution magnetohydrodynamics simulation of interchange reconnection in a pseudostreamer topology by P. F. Wyper et al. A remarkable similarity between the simulation-derived images and the observations was found. We conjecture that these Metis observations may represent the upper ends of the spatial and energy scales of the interchange reconnection process that has been proposed recently as the origin of the Alfvénic solar wind.

*Unified Astronomy Thesaurus concepts:* [Solar wind \(1534\)](#)

*Materials only available in the [online version of record](#): animations*

## 1. Introduction

The Metis coronagraph (E. Antonucci et al. 2020; S. Fineschi et al. 2020), on board the ESA's Solar Orbiter mission (D. Müller et al. 2020), represents a major advancement in our ability to observe and understand the solar corona. Launched in

2020 February, the Solar Orbiter aims to study the Sun up close, and Metis plays a crucial role in this mission by providing high-resolution images of the outer corona in both visible and ultraviolet light. By occulting the bright solar disk, Metis enables scientists to observe the faint corona, unveiling intricate details of the Sun's outer atmosphere and its dynamic behavior (M. Romoli et al. 2021). The Metis coronagraph's advanced imaging capabilities and its integration with the Solar Orbiter's suite of instruments are enhancing our understanding of the coupling between the corona and solar wind, offering



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valuable data that support both theoretical models and practical space-weather applications.

Metis measures the brightness of the corona in two wavelength bands: broadband visible light (VL) in the range 580–640 nm and narrowband ultraviolet radiation in the range  $121.6 \pm 10$  nm. The instrument is capable of observing at cadences that can be as high as one image per second, thus allowing the study of the structure and dynamics of the coronal plasma and magnetic fields, including the evolution of coronal mass ejections (CMEs), waves, and other transient events.

Two important developments in recent years in the observation and theory of corona–solar wind coupling have made the Metis capabilities especially valuable. First, there is a growing consensus that the so-called Alfvénic wind from coronal holes is due to ubiquitous jetting activity that is observed at the base of the corona. The jets are presumed to result from the interchange reconnection between the closed flux of small-parasitic-polarity regions and the surrounding coronal-hole open field (S. D. Bale et al. 2023; N. E. Raouafi et al. 2023). Supporting evidence for this hypothesis comes from in situ measurements by the Parker Solar Probe (PSP), which demonstrate that the bulk of the solar wind exhibits fluctuations that are Alfvénic, in that the magnetic and velocity perturbations are tightly correlated (P. Thepthong et al. 2024). These fluctuations are believed to originate in the low corona and are a natural consequence of the interchange reconnection responsible for the jets. It should be noted that the idea of a reconnection-driven origin for the wind is not new. E. N. Parker (1992) and W. I. Axford & J. F. McKenzie (1992) independently proposed that the heating and acceleration of the wind is due to reconnection between small-scale closed-field regions and background open flux. The recent high-resolution observations of the corona and photosphere strongly support that hypothesis (S. D. Bale et al. 2023; N. E. Raouafi et al. 2023; L. P. Chitta et al. 2024). The ubiquitous jets drive both mass motions and Alfvén waves onto open field lines, but due to their small scale, only the mass flow has been imaged directly. We emphasize that these ubiquitous jets are well below the Metis field of view (FOV); consequently, Metis cannot test this model directly. Metis, however, can look for evidence of outflows from large-scale interchange reconnection events.

The second important development is high-resolution observations have also revealed that coronal jets are due to the eruption of a filament, just like large-scale eruptions such as CMEs (A. C. Sterling et al. 2015). It appears that in all cases, the free energy to power eruptive events builds up at the base of the low corona, in the form of the highly sheared magnetic field of a filament channel overlying a polarity inversion line (PIL). In the case of coronal hole jets and pseudostreamer eruptions, the pre-eruption magnetic topology is that of a closed-field embedded bipole surrounded by open flux (e.g., E. Pariat et al. 2009). This topology is physically identical to that of the breakout model for CME onset (S. K. Antiochos et al. 1999), in that it contains a dome-like separatrix surface in the corona with a null point. Reconnection at this coronal null point and separatrix allows the filament-channel field to expand upward, causing a current sheet to form below the rising filament, which leads to strong flare reconnection there and eventually results in a jet-like eruption through the null points. We have argued and demonstrated with numerical simulations that this basic

scenario is a universal mechanism for solar eruptions (P. F. Wyper et al. 2017).

The key conclusion from these observational and theoretical developments is that we can probe the basic process that has been proposed for the origin of the Alfvénic solar wind by studying pseudostreamer eruptions, which are large-scale and can be well observed by Metis and other instruments. This is the primary motivation for the work presented below. Our results capture these processes of interchange-driven plasma and Alfvén wave injection from the corona to the heliosphere.

The discussion above emphasizes the central role of erupting filament channels in powering the interchange reconnection that drives the wind. Although the pre-eruptive magnetic structure of the filament channels is still under debate, there is widespread consensus that the final erupting structure that escapes from the corona as a CME, or that reconnects through a null point as a jet, is a twisted flux rope, due to the action of flare reconnection (E. R. Priest & T. G. Forbes 2002). Note that by “flux rope” we refer to a magnetic flux tube in which the field lines spiral around a central axis. Flux ropes are the fundamental structures that appear high in the corona during eruptive events, and in situ measurements of interplanetary CMEs generally show a flux-rope structure with significant twist/helicity (e.g., S. Dasso et al. 2005). Metis observations have already proven effective for revealing the presence of eruptive flux ropes in the corona, providing evidence of their footpoints and helping to map their 3D structure. Flux ropes are known to play a critical role in the dynamics of the solar atmosphere (S. E. Gibson 2018) and the acceleration of CMEs (D. F. Webb & T. A. Howard 2012) and are often observed as filamentary structures in the solar corona (R. Liu 2020).

In the context of CMEs, flux ropes are crucial, because they carry away the magnetic free energy and helicity (P. Romano et al. 2003, 2014). The flux rope usually forms the bright core of a CME, and in the standard MHKSP model for eruptive flares (H. Carmichael 1964; P. A. Sturrock 1966; T. Hirayama 1974; R. A. Kopp & G. W. Pneuman 1976), it is located above the flare current sheet in the traversal direction to the underlying cusp structure and is created by the flare reconnection (E. R. Priest & T. G. Forbes 2002). When a flux rope erupts from inside a pseudostreamer, however, it generally undergoes interchange reconnection through the null point with the surrounding open flux, so the resulting ejection is a jet or a narrow fan-like CME (Y. M. Wang & P. Hess 2018; R. A. Kopp & G. W. Pneuman 2021; P. F. Wyper et al. 2021). The free energy and helicity contained in the flux rope are converted to mass acceleration and an Alfvén wave flux by interchange reconnection, as required by the recent solar wind models (S. D. Bale et al. 2023; N. E. Raouafi et al. 2023).

Recent observations and simulations have highlighted the presence of helical and twisting structures in the solar corona, contributing to a better understanding of coronal dynamics and their relation to the solar wind. In particular, coronal jets and jetlets have been extensively studied as potential drivers of Alfvénic disturbances propagating into the heliosphere. For instance, A. C. Sterling et al. (2020) discuss the possible evolution of minifilament-eruption-driving coronal jets into magnetic-twist-wave “switchbacks” observed by the PSP (see also R. Biondo et al. 2023). Their study emphasizes that the twisting motions observed in jet spires may propagate outward, forming magnetic structures detectable as transient magnetic field reversals in situ.

**Table 1**

Summary of Metis Data Sets Taken on 2022 October 12 and Analyzed in This Work: Type of Acquisition (“Polarized Brightness”—pB or “Total Brightness”—tB), Time Intervals Covered, and Main Parameters of the Acquisition Session (Exposure Time, Cadence, and Detector Binning)

Index	Type	Session No.	No. of Images	Start Time	End Time	Cadence (s)	Binning
1	VL-pB	228501	19	03:30:01	08:48:23	960	$2 \times 2$
2	VL-tB	228504	120	09:30:31	10:12:03	21	$2 \times 2$
3	VL-tB	228505	299	10:15:01	20:18:45	121	$2 \times 2$
4	VL-pB	228506	36	20:25:17	23:24:13	300	$1 \times 1$

**Note.** Each data set is identified by a session number that uniquely identifies in the Metis database the images of that data set. Times are onboard UTC.

Furthermore, evidence from observational campaigns, including total- and polarized-brightness images, supports the interpretation that helical structures are a natural consequence of the reconnection process in the corona. For example, S. R. Habbal et al. (2014) provide insights into the polarimetric signatures of helical structures and their evolution in the corona, revealing key characteristics of these configurations at various altitudes.

In addition, white-light images from coronagraphs have shown the presence of propagating helical features consistent with theoretical predictions of magnetic flux-rope dynamics in CMEs and jets. For example, M. Druckmüller et al. (2014) demonstrated the use of advanced image processing techniques to uncover fine-scale helical patterns in white-light observations, revealing the complexity of the coronal magnetic field’s interaction with the solar wind. Similarly, R. A. Harrison et al. (2001) reported observations of twisted structures in the extended corona, interpreted as large-scale helical flux ropes forming in the aftermath of solar eruptions. These observations highlight that such helical configurations are not limited to small-scale jet-like structures but can also be found in larger-scale coronal phenomena, further supporting the role of magnetic reconnection in shaping the solar wind’s properties.

In this work, we analyze high-cadence Metis observations taken in both unpolarized and polarized VL, contextualized by Extreme Ultraviolet Imager (EUI) imaging, providing new insights into the nature of the corona–solar wind coupling (see V. Andretta et al. 2021; A. Bemporad et al. 2022; G. Russano et al. 2024). These observations captured a jet-like CME on 2022 October 12, associated with the eruption of a polar crown prominence from beneath a pseudostreamer. We interpret the event in the framework of the breakout jet model (e.g., P. F. Wyper et al. 2018, 2021; R. A. Kopp & G. W. Pneuman 2021) and investigate the post-eruption relaxation phase, which exhibits evolving filamentary structures. To further understand these features, we compare Metis observations with synthetic white-light images derived from the high-resolution magneto-hydrodynamic (MHD) simulation of slowly driven interchange reconnection performed by P. F. Wyper et al. (2022), revealing similarities that offer new insights into coronal dynamics and the processes believed to drive the solar wind.

The paper is organized as follows. The next section describes the data used and gives an overview of the observations. In Section 3, we give our theoretical interpretation of the outflows observed by Metis. Section 4 details the methods employed to measure the inclination and pitch of the helical structures and the obtained results. Section 5 then outlines the comparison with the MHD simulation, while the final section presents the conclusions.

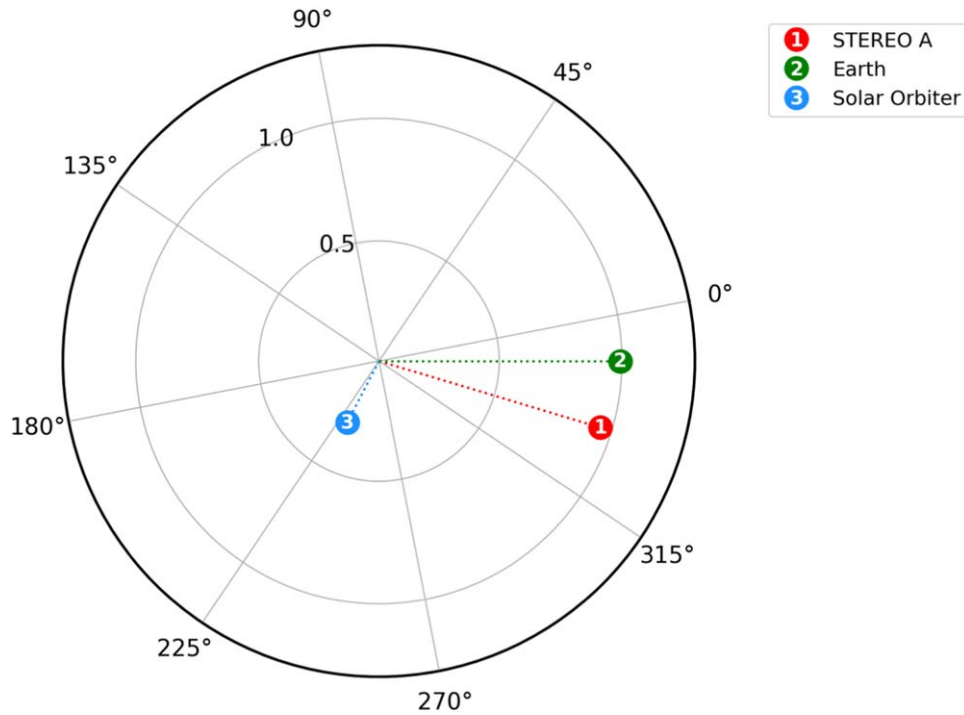
## 2. The Event of 2022 October 12 Observed by Metis

For our analysis, we used Metis coronagraphic images in VL acquired on 2022 October 12, from 03:30:01 UT to 23:24:13 UT. The data were acquired both in the polarized-brightness observing mode (VL-pB) and in the so-called “total-brightness” mode (VL-tB). The former consists of the acquisition of four interlaced polarized VL images that are then combined to obtain a map of the corona in polarized brightness (*pB*), while the latter acquires the total (unpolarized) coronal brightness through a special hardware mode that switches the polarization angle in the middle of the detector integration time (E. Antonucci et al. 2020). Table 1 lists the Metis VL observations on that date with their main acquisition parameters. On that date, the Solar Orbiter was a distance of 0.293 au from the Sun, pointing to a position on the solar surface centered at a Carrington longitude and latitude of about  $230^\circ$  and about  $-3^\circ$ , respectively. At that distance, the angular pixel scale of the detector of the Metis VL channel is  $10''.7$  (V. Andretta et al. 2021), corresponding to about 27 Mm on the Sun (54 Mm if a  $2 \times 2$  binning is applied to the data). Figure 1 shows the position of the Solar Orbiter relative to Earth: as we can see, they are almost in quadrature, which means that the western corona in the Metis images lies along the Earth line of sight.

The highest-cadence observations were obtained in the total-brightness mode, with cadences of about 20 s from 9:30 to 10:12 UT and about 2 minutes from 10:15 to 20:19 UT. The data have been calibrated, as explained in M. Romoli et al. (2021) and Y. De Leo et al. (2023). Further processing steps were applied to the image sequences to emphasize the most rapidly varying features in the FOV. In particular, for this work, we chose to show running-difference images normalized by the average radial coronal brightness profiles. In order to increase the signal-to-noise ratio, we also applied a pixel-by-pixel temporal average to the highest-cadence VL-tB acquisition (session no. 228504), effectively reducing the cadence to one image every 140 s in that case.

Figure 2 reports four *pB* Metis images processed using the normalized running-difference technique between two successive images, acquired with a 16 minutes cadence. This sequence reveals the propagation of a jet-like CME at approximately a  $35^\circ$  polar angle (measured from solar north in a counterclockwise direction). Initially (the top left panel of Figure 2), the corona in the northeastern quadrant appears dominated by radial structures. Starting from 05:38 UT, a dark structure (see the white arrow in the top right panel of Figure 2), likely the CME front, becomes visible in the running-difference maps, flanked on both sides by two bright features extending to the upper edge of the FOV (see the black arrows in the top right panel of Figure 2). Note that the





**Figure 1.** Positions of the Solar Orbiter and STEREO A with respect to Earth on 2022 October 12 6:30 UT, in heliocentric Earth equatorial coordinates. The image was created using Solar-MACH (<https://solar-mach.github.io/>).

easternmost bright radial feature was present before the event, brightens during the eruptions, and, as will be shown below, persists well after the CME. We identify this feature as a pseudostreamer stalk, which corresponds topologically to either an open spine or fan emanating from a coronal null point (e.g., P. F. Wyper et al. 2021). In response to the CME eruption, the bright features seem to gradually move apart (as seen in the bottom panels of Figure 2), with a separation velocity of about  $100 \text{ km s}^{-1}$  at a distance of approximately 1.5 solar radii. Interestingly, by 06:26 UT, a broad dimming between the bright features has developed (see the white arrows in the bottom right panel of Figure 2). Similar dimmings have been previously reported in jet-like CMEs (e.g., Y. M. Wang & P. Hess 2018; R. A. Kopp & G. W. Pneuman 2021).

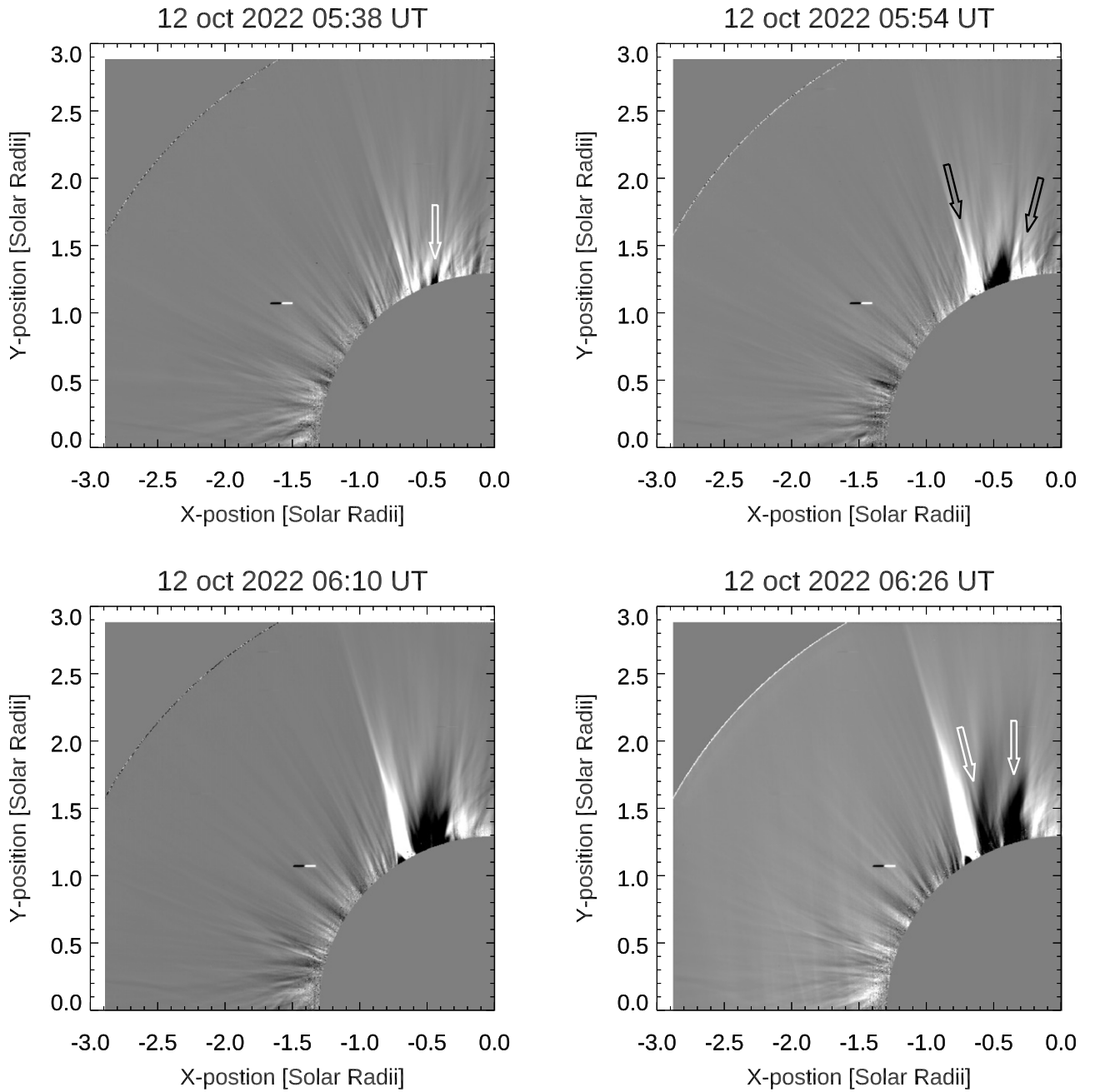
In order to highlight the substructures within the CME and surrounding corona, we applied a normalized base-difference technique. This technique involves computing a base image from the series of images by taking the pixel-by-pixel first percentile. The base image is then subtracted from each image and used to compute a mean radial profile, by averaging its values in the azimuthal direction. The image containing the radial profile is then used to normalize the images, enhancing the contrast at higher radial distances. Subsequently, we applied the standard running-difference technique between two consecutive images (V. Andretta et al. 2025).

Applying the normalized base-difference technique between images with a cadence of 2 minutes and 1 s, at 10:15 UT, we observe that the eastern feature, which we believe corresponds to the pseudostreamer stalk, exhibits a markedly helical configuration along its entire radial extent within the Metis FOV, from about 1.5 to approximately 3 solar radii ( $R_s$ ; see Figure 3). This configuration persists for more than 3.5 hr from the beginning of the total-brightness image sequence (see Figure 14). We also note faint striations in the western feature and in the gap between them, but the stalk is unique in its

pronounced brightness, striking helical dynamics, and persistence. These observations show exactly the type of structure and dynamics that have been postulated by N. E. Raouafi et al. (2023) and others as the origin of the Alfvénic solar wind, but on a much larger scale. We discuss this point in more detail directly below.

The observed narrow CME and the continued helical structure appear to be the result of an eruptive polar crown prominence evident in the 174 Å images acquired by the EU1 (P. Rochus et al. 2020), also on board the Solar Orbiter, starting at 4:00 UT. These images show the eruption of a prominence at the same latitude near the north solar pole on the eastern limb (the top left panel of Figure 4). Applying the running-difference technique to the EU1 images, we highlight the twist of the erupting prominence, manifesting a helical configuration during its rising phase. In particular, at 6:10 UT, the eruptive prominence exhibits an X-shaped structure, with the narrowest part located at a height of approximately  $1.1 R_s$  (the top right panel of Figure 4). Ahead of the erupting prominence, a linear feature also becomes visible around 5:10 UT and moves upward and eastward, disappearing around 7:10 UT (see the supplemental animations of Figures 14 and 15). After this time, the erupting prominence loses its coherence, and it fades from view by 8:00 UT. Finally, over the next 12 hr, faint cusped loops form in addition to two faint ribbons (the bottom left and right panels of Figure 4).

Taking into account the high latitude from which the CME originates, we might expect that despite the relative positions of the Solar Orbiter and STEREO A relative to Earth, the CME and prominence eruption might also be detected. Indeed, the prominence eruption appears above the limb in the STEREO EUVI from around 6:10 UT onward (see Figure 5) and follows a similar evolution of erupting and losing its coherence around 7:00 UT. At 7:32 UT, a faint jet-like CME appears in the Large Angle and Spectrometric Coronagraph FOV.



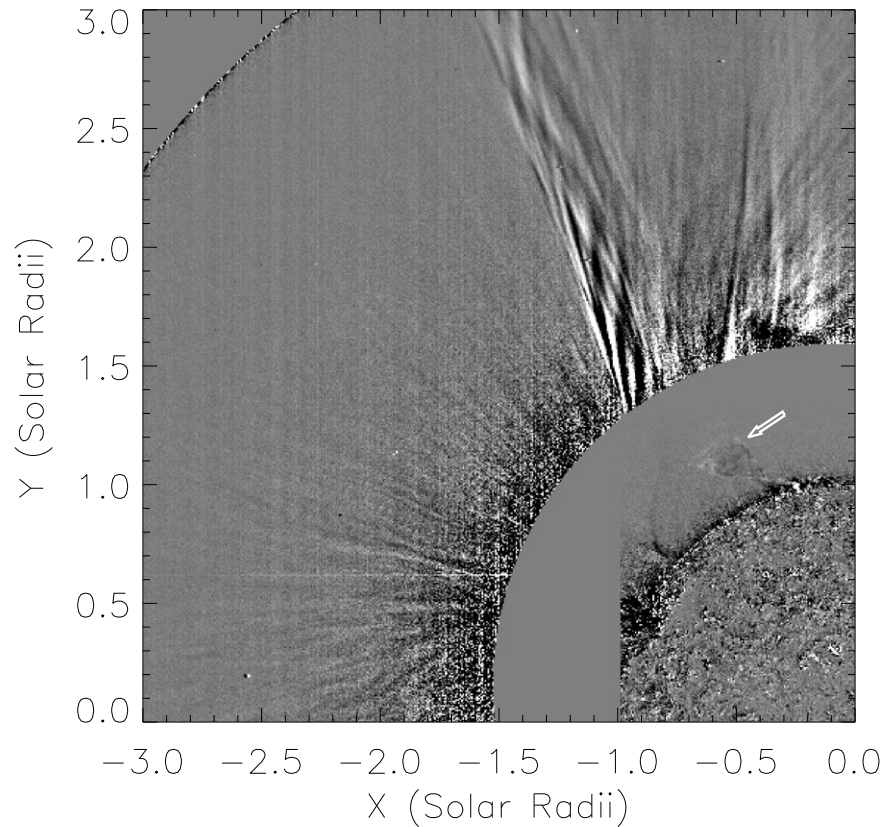
**Figure 2.** Sequence of the normalized running-difference maps depicting the propagation of the CME front (white arrows) and the evolution of its two legs (black arrows). The time indicated on each map refers to the image from which the previous one, acquired 16 minutes earlier, has been subtracted.

### 3. Theoretical Interpretation

In order to determine the magnetic field configuration of the outer corona in the observed region, we used the Magnetohydrodynamics outside A Sphere (MAS) method. MAS is a 3D MHD model (J. A. Linker et al. 1999; Z. Mikić et al. 1999, 2018), which uses the magnetic flux distribution at the base of the computational domain for boundary conditions, includes thermal conduction along field lines, radiative losses, and coronal heating, and integrates the time-dependent equations in spherical coordinates until the configuration reaches a steady state (e.g., L. Abbo et al. 2015, which employed the simpler polytropic approximation). The obtained coronal magnetic-field-line map extrapolated through the model for the north-eastern quadrant up to 3 solar radii is shown in Figure 6, where

we can identify a pseudostreamer at the latitude corresponding to the location of the polar crown prominence.

The EU images show that the polar crown prominence resides initially under the northern arch or lobe of this pseudostreamer. The eruption of the prominence follows the typical evolution of a jet-like pseudostreamer eruption, which is itself very similar to a large-scale version of a coronal jet involving a small-scale filament. Both kinds of eruption have been shown to be well explained by the breakout jet model (e.g., P. F. Wyper et al. 2018). Here, we interpret the linear feature that forms around 5:20 UT ahead of the erupting prominence as a breakout current layer formed at the apparent null point of the pseudostreamer (which is likely a chain of null points and separators; e.g., V. S. Titov et al. 2012). The helical structure within the erupting prominence reflects the expected



**Figure 3.** Composite image combining an EUV map at 174 Å taken at 6:10 UT and a Metis map captured at 10:17 UT. The normalized running-difference technique has been applied to both maps. The arrow points to the eruptive polar crown prominence depicted in Figure 4.

flux-rope structure that could be pre-existing or formed during the eruption, with the X-shaped region showing the likely position of the flare reconnection beneath the flux rope. The loss of coherence and disappearance of the erupting prominence are consistent with the reconnection of the prominence flux rope at the breakout current sheet around 7:00 UT. As in previous breakout jet simulation studies (e.g., P. F. Wyper et al. 2018, 2021), this is expected to transfer some of the twist of the flux rope to the open field, launching helical outflows along the open field consistent with the initial jet-like CME and helical motions observed by Metis.

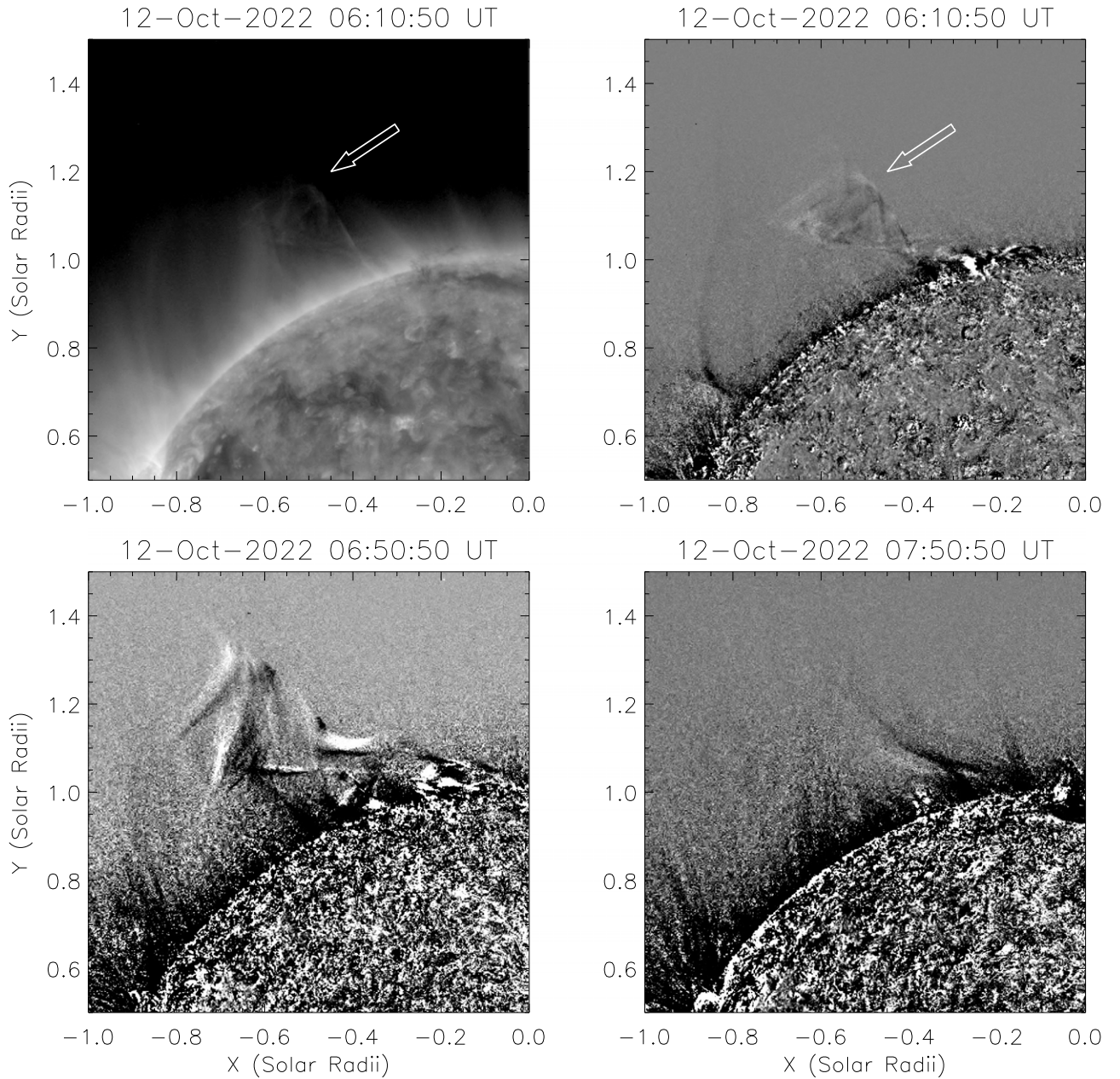
The observations described above of the erupting prominence and jet-like CME are fairly typical and similar to those of many other events. The striking new features of the Metis observations are the helical motions along the radial stalk that continue to enter the FOV well beyond the initial opening of the flux rope and disappearance of the prominence in the EUV around 7:00 UT and the outward propagation of the jet-like CME. These helical motions are still clearly discernible until at least 11:30 UT, far into the relaxation phase of the pseudostreamer after the eruption. We claim that these motions are a result of the “rebound” reconnection of the closed flux system, which is a natural—but until now—unobserved consequence of the evolution expected for a pseudostreamer CME or coronal-hole jet.

We show the basic scenario schematically in Figure 7. For ease of viewing, the system is shown in rough 2D, in that there are two PILs (dashed white lines) on the photosphere and two separate closed arcades. In reality, the PILs loop around and form one elliptical PIL, but the schematic captures the basic physics of the system. The breakout and flare reconnection are

shown in red and blue, respectively, in panel (a) during the eruption phase. Panel (b) shows that when the flux rope undergoes reconnection through the null point, part of it connects to the open field, which launches the initial jet-like CME, while the rest connects to the closed field above the left PIL. Note that in this initial reconnection, only half or so of the shear/twist in the erupting flux rope escapes out into the heliosphere as a CME, with the rest staying inside the pseudostreamer as newly closed flux overlying the nonerupting portion of the PIL. This result is a basic property of all interchange reconnection; only the stress on the leg of the initially closed field line that becomes open is released into the heliosphere. As a result, the closed-field region on the left in Figure 7 now has an excess of magnetic flux and shear/twist. Once the filament flux rope exits the pseudostreamer, this closed flux “rebounds” back toward the right, as in panel (b), leading to a second round of reconnection, now between the left closed system and the open flux on the right. This rebound reconnection transfers further twist to the open field, leading to pronounced helical outflows in the relaxation phase. Such double reconnection of the flux rope in breakout simulations has been noted before (see the silver field lines in Figure 7 of P. F. Wyper et al. 2021, for example), but since the simulations were focused on the eruptive phase, this relaxation phase was not explored in detail. Note also that the rebound interchange reconnection again leaves some magnetic shear/twist in the system, so there could be continued back-and-forth rounds of successively less energetic interchange reconnection.

We show schematically in Figure 8 our interpretation of the event within this framework. The broad region of the CME in the Metis FOV is explained by the jet-like outflows from the

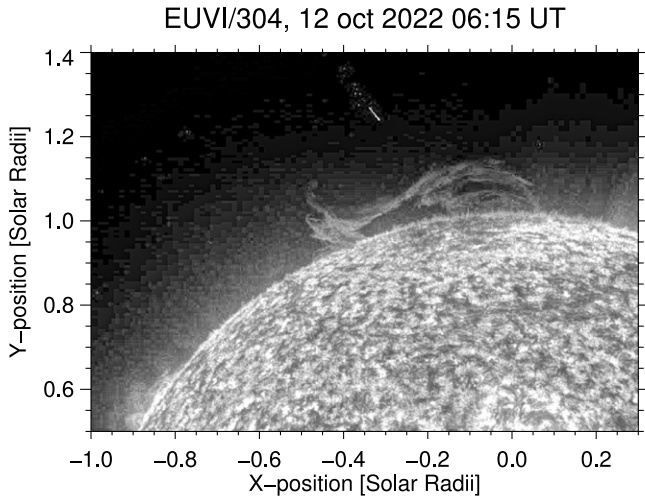




**Figure 4.** In the top left panel, an EUV map at  $174 \text{ \AA}$  is shown, captured on 2024 October 12, at 6:10 UT. The top right panel displays the same FOV using a map generated with the running-difference technique. Arrows indicate the eruptive polar crown prominence, with its helical configuration clearly visible during the rising phase in the running-difference map. The bottom left and bottom right panels present running-difference maps taken at 6:50 UT and 7:50 UT, respectively.

eruption following the fan-like open spine of the pseudostreamer. The bright easternmost feature in Figure 2 corresponds to where the pseudostreamer open spine or fan aligns with the line of sight, while the on-disk extensions of the fan-like outflows lead to the relatively broad region of helical and striated outflows (panels (a) to (c)). The helical outflows in the late phase of the event, located primarily on the easternmost edge, occur as further twist is released by the rebound reconnection of the closed field above the southernmost PIL, around the same time that the cusp-shaped loops and faint flare ribbons form. Based on this timeline, we conclude that the Metis observations show outflows and helical motions occurring as a result of long-duration interchange reconnection in the aftermath of the polar crown prominence eruption.

The key point of the rebound or relaxation reconnection is that it is completely generic. The CME itself may have unique properties, because it involves flare reconnection and most likely the large deformation of the pseudostreamer topology, but the late-phase evolution corresponds to the interchange reconnection that any parasitic-polarity region will inevitably have with the surrounding open flux. Even if a bipolar region emerges in a purely potential state, which is highly unlikely, the constant photospheric convective motions are bound to drive interchange reconnection. Both the PSP (N. E. Raouafi et al. 2023) and very recent Solar Orbiter observations (L. P. Chitta et al. 2024) imply that this type of interchange reconnection as observed by Metis, but on a much smaller spatial and energy scale, is the origin of the Alfvénic solar wind.



**Figure 5.** STEREO EUVI image at 304 Å of the polar crown prominence taken at 6:15 UT. The image is processed with Multiscale Gaussian Normalization (H. Morgan & M. Druckmüller 2014).

#### 4. Measurements of the Helix Parameters

By using total-brightness images of the extended corona and applying the running-difference technique, we are able to visualize more clearly the helical structure on the easternmost extent of the outflows. Specifically, the helical structure emerges from the alternating bright and dark filamentary features that display a greater inclination relative to the longitudinal axis of the flux tube that forms the pseudostreamer stalk. Here, we define “inclination” as the angle between these filamentary features and the radial direction originating from the solar disk passing through the midpoint of the intersecting leg of the flux tube and the solar disk. The inclination provides a measure of how much these features deviate from the radial direction. The dark features emerging from the running-difference maps are analyzed to determine their inclination.

In both panels of Figure 9, two running-difference maps are shown as examples, with the dark features indicated by red segments, with the inclinations being measured and displayed in the top left of each panel. The number of identifiable features varies among the frames, and specific features and their evolution from one frame to the next cannot always be consistently tracked. Nonetheless, a consistent trend is observed among the identified features in individual frames, where features located at higher polar angles and higher radial distances exhibit lower inclinations, i.e., closer to the radial direction. In the two frames of Figure 9, the inclination changes from approximately  $30^\circ$  to  $16^\circ$ ; a similar trend is visible in the subsequent frames.

Subsequently, the height of the dark features—referred to as the helix pitch—was measured by determining the distance of the intersection points between these features and the radial direction passing through the midpoint of the intersecting leg of the flux tube and the solar disk. Considering that features at higher coronal polar angles and greater radial distances do not always span a sufficient length to intersect the radial direction defined by the midpoint of the leg, this estimation of the height was performed only for some features in each frame.

The temporal evolution of the inclination and helix pitch, utilizing the first 20 available frames, is depicted in the top and bottom panels of Figure 10, respectively. It is found that the average inclination of the helical features varies from  $31^\circ$  to  $6^\circ$ .

The span of the radial distance of the intersection points between the dark features and the radial direction remains relatively constant over time, varying between  $1.5$  and  $2.0 R_s$ .

Considering a longer time sequence, from the beginning of the available total-brightness Metis sequence until the disappearance of the helical structure (around the hundredth Metis frame, corresponding to 13:22 UT), we note that the inclination of the dark features tends to become more spread out, as highlighted by the standard deviation of the mean, shown in gray in the top panel of Figure 11. Additionally, there is an increase in the height of the higher intercepts between the dark features and the radial direction, especially toward the end of the sequence, suggestive of an overall reduction in twist within the flux tube toward the end of the observation period (the bottom panel of Figure 11).

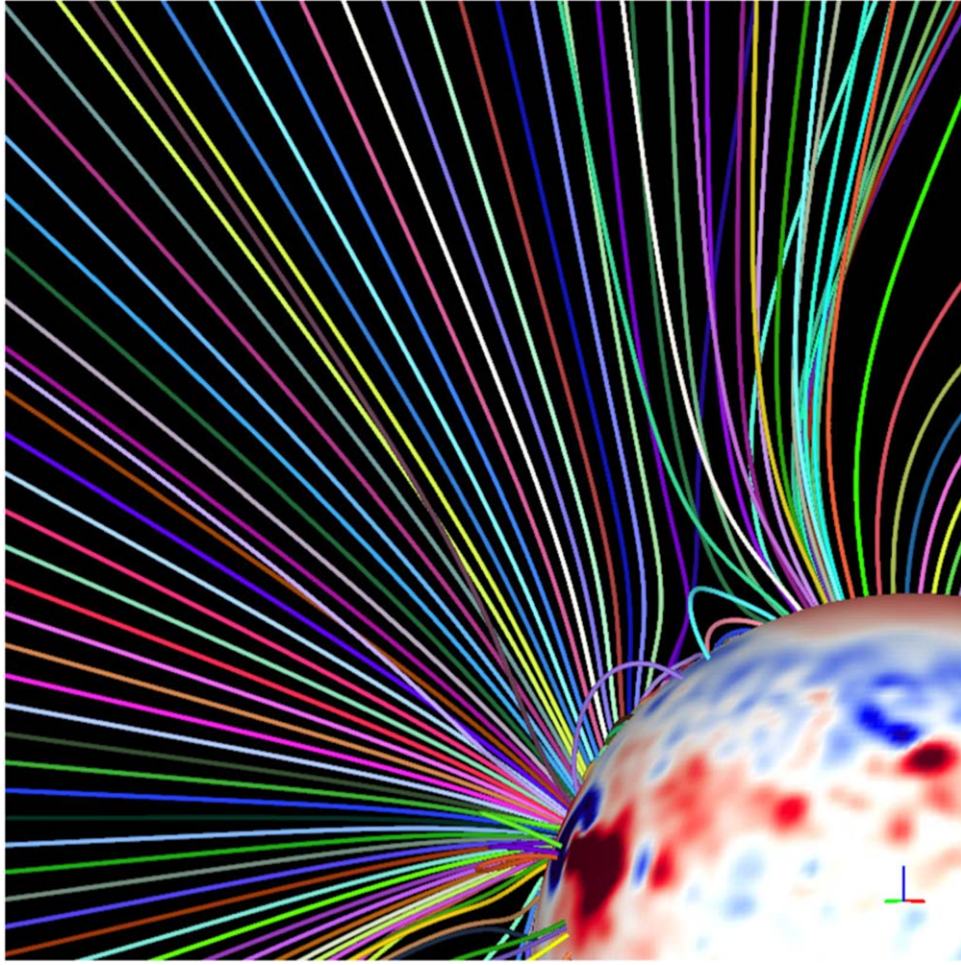
Observing the plot in the bottom panel of Figure 11 more closely, we note that some points, likely belonging to the same feature visible in several consecutive frames, display their rising process by the continuous monotonic increase of their height: from these sequences of points, we can estimate an ascent velocity of approximately  $0.5 R_s$  per hour, corresponding approximately to  $100 \text{ km s}^{-1}$ .

#### 5. Comparison with an MHD Simulation of Interchange Reconnection

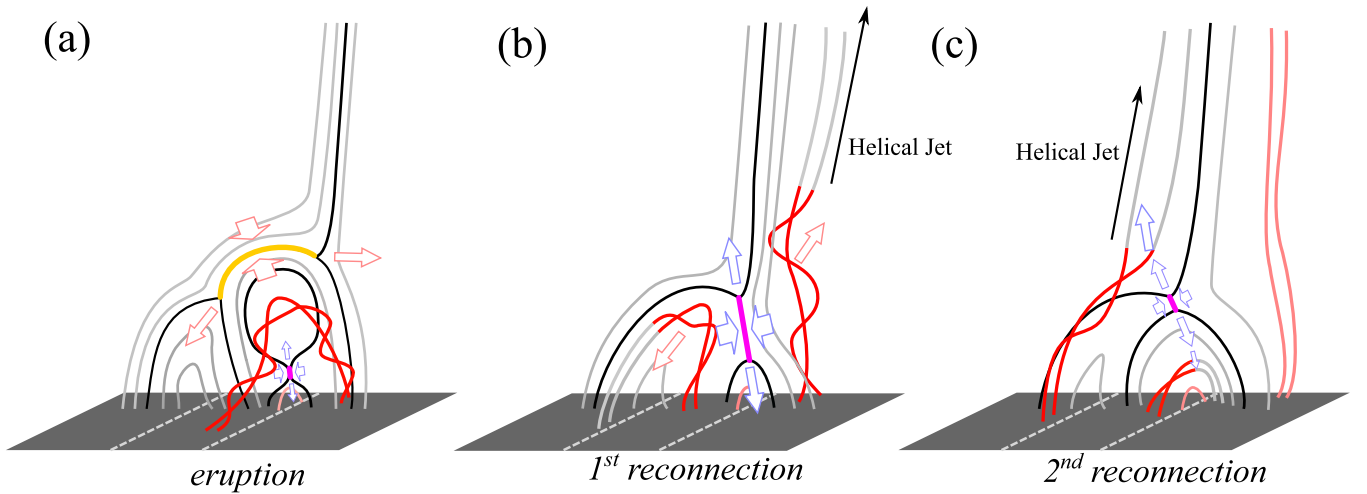
Considering our interpretation that Metis was observing outflows and helical motions as a result of sustained interchange reconnection, we decided to investigate the similarities between the Metis observations and the results of the high-resolution simulation of interchange reconnection conducted by P. F. Wyper et al. (2022). The simulation domain was the volume between between  $1$  and  $20 R_s$  (although the outflows were only resolved as far as  $4.5 R_s$ ), and the system was initialized with a monopolar ambient magnetic field combined with 16 subsurface magnetic dipoles, resulting in a 3D null-point topology similar to the pseudostreamer in this event. The magnetic null point in the simulation was at a height of about  $1.25 R_s$ . A fully ionized plasma was assumed, with a constant and uniform temperature of  $1 \text{ MK}$  throughout the computational volume.

Interchange reconnection was initiated in this simulation by applying twisting motions at the bottom boundary corresponding to the photosphere. Note that the goal of the simulation was to investigate the fundamental process of interchange reconnection, not CMEs and/or prominence eruptions, so the motions were large-scale and slow and simply added a global stress to the closed flux system. As a result of the added magnetic stress, the closed flux expanded upward, deforming the null point and separatrix into a current sheet. An important feature of the simulation is that it used adaptive mesh refinement to achieve the highest possible resolution of the current sheet and ensuing reconnection dynamics. Being highly resolved, the current sheet fragmented due to the plasmoid instability, and plasmoids with enhanced density repeatedly formed and were ejected from the current layer. Plasmoid formation is expected to be a general property of all high-Lundquist-number reconnection (A. Bhattacharjee et al. 2009). As discussed in detail in P. F. Wyper et al. (2022), the field line twist within the plasmoids propagated away as torsional Alfvén waves. But of most relevance here is that the overdense plasma of the plasmoids then followed behind these waves as field-aligned dense outflows. Furthermore, due to the added





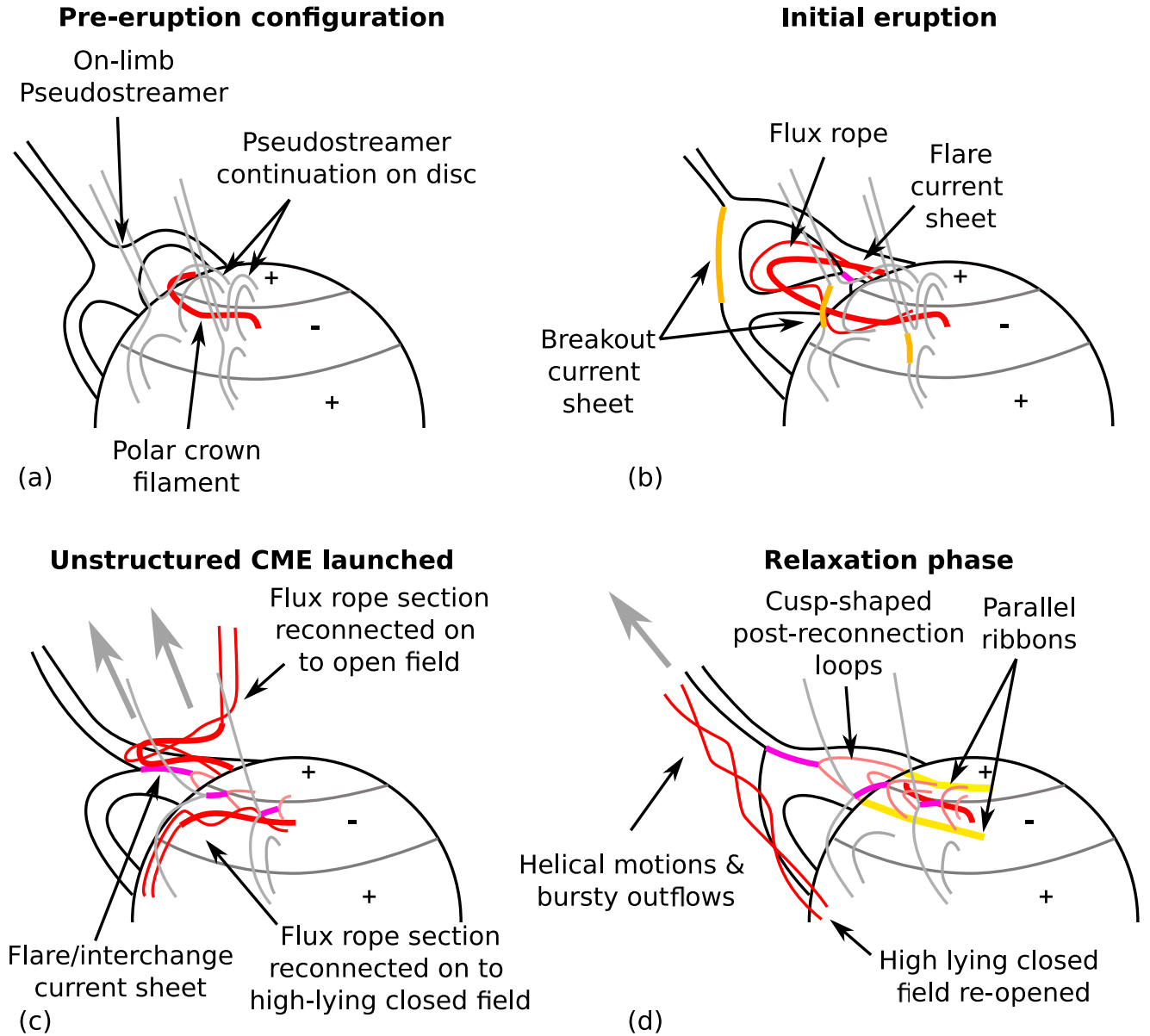
**Figure 6.** Extrapolations of the coronal magnetic field on 2022 October 12. We have identified a pseudostreamer at the northeast limb at the latitude corresponding to the location of the polar crown prominence. The FOV extends up to 3 solar radii.



**Figure 7.** Schematic showing a scenario for how helical jet-like motions can arise in the relaxation phase of the breakout jet model. (a) The initial eruption. The breakout and flare current sheets are shown in yellow and magenta, respectively. The arrows indicate the direction of the reconnection and plasma movement in each. The dashed lines show the PILs. (b) The flux rope (red) is split when it reaches the breakout current layer and is reconnected. (c) Interchange reconnection in the relaxation phase transfers magnetic flux back over the right PIL, transferring more closed-field twist to the open field line.

magnetic stress, the closed field had a shear (nonpotential) component. As a result, when this closed field reconnected with the open field, the newly formed open field lines had a large-

scale kink, i.e., a nonlinear Alfvén wave, which straightened out over time. Combined with the repeated launching of dense outflows along these field lines by the plasmoid ejections, this

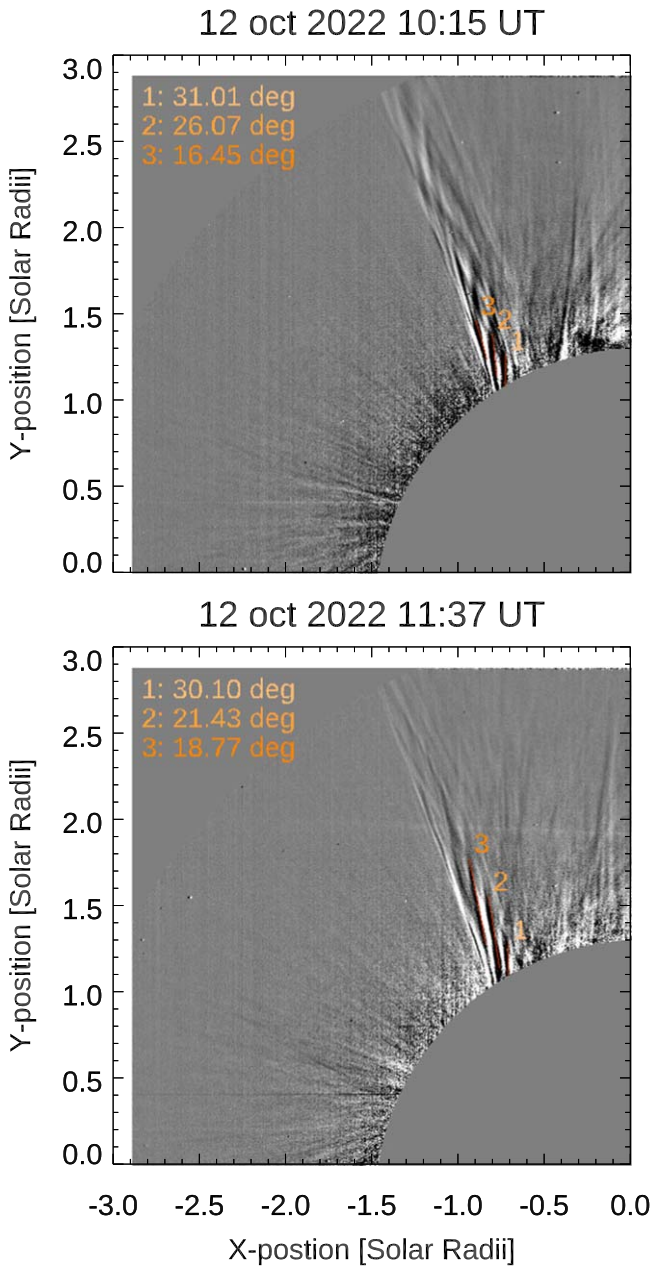


**Figure 8.** Schematic of the hypothesized magnetic field evolution during the event. See Section 3 for details.

led to a curtain of dense filamentary structures that gradually straightened out.

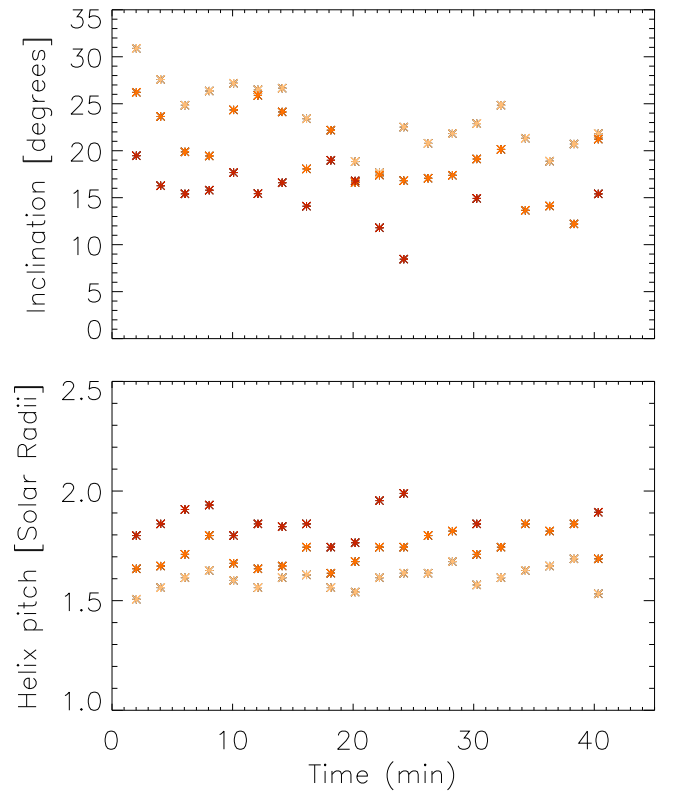
When viewed in synthetic white-light images derived from the simulation density (see B. J. Lynch et al. 2025 for details), these dense features appear very similar to the ones observed by Metis. Therefore, we conducted an analysis similar to the one performed on the Metis data set, measuring the inclination of the inclined features and the inferred height on the running-difference maps derived from the synthetic white-light images created from the simulation. Also in this case, dark features were examined, assuming analogous results for the bright features. As observed in the total-brightness images, the simulations also exhibit higher inclinations of the lower features located at smaller polar angles. Synthetic images allowed for the segmentation of a greater number of features, taking into account that the outflows are resolved out to  $4.5 R_s$  (for example, five and six dark features are identified in the frames shown in the top and bottom panels of Figure 12, respectively). In the simulation, the range of inclinations

appears broader, spanning from about  $28^\circ$  to  $-4^\circ$  (the negative inclinations correspond to features forming an angle opposite to the radial direction). However, it is important to acknowledge the limitations inherent in the use of time-differencing techniques to highlight these fine-scale dynamic features. While running-difference maps are essential for revealing rapid changes in brightness and detecting propagating structures, they can introduce artifacts or obscure slow-evolving features that may not be clearly visible in individual frames. Additionally, the choice of cadence can influence the visibility and apparent evolution of features, potentially affecting the measured inclination and spatial distribution of the dark and bright structures, although it is noteworthy that the helical features are also visible in the running-difference maps obtained using the data set with a cadence of 21 s (the third row of Table 1). Despite these limitations, the consistency observed between the Metis observations and the synthetic images from the MHD simulation reinforces the robustness of our interpretation.



**Figure 9.** Two Metis maps of the northeastern portion of the corona obtained by the running-difference technique. In the top left of each panel, the angle between each structure and the radial direction is indicated. The numbers and the red levels distinguish the different features.

To make a more quantitative comparison between the Metis observations and the simulation, we considered the distribution of the radial distance of the features as a function of their inclination (the top and bottom panels of Figure 13, respectively). In the case of the observations, the radial distance was taken as the intersection of the radial direction with the features, while for the simulation, the midpoint of the features was considered. Although the helical features are located at different altitudes in the observations and the simulation, likely due to the boundary conditions of the latter not being perfectly comparable with the observations, both distributions show a similar anticorrelation between the two quantities, i.e., as the radial distance increases, the inclination decreases. Notably, in the case of the simulation, whose



**Figure 10.** Temporal evolution of the inclination and helix pitch of the helicoidal structures, as observed by Metis.  $t = 0$  corresponds to 10:15:01 UT on 2022 October 12. The varying shades of red represent features at progressively greater radial distances, with darker shades indicating higher altitudes.

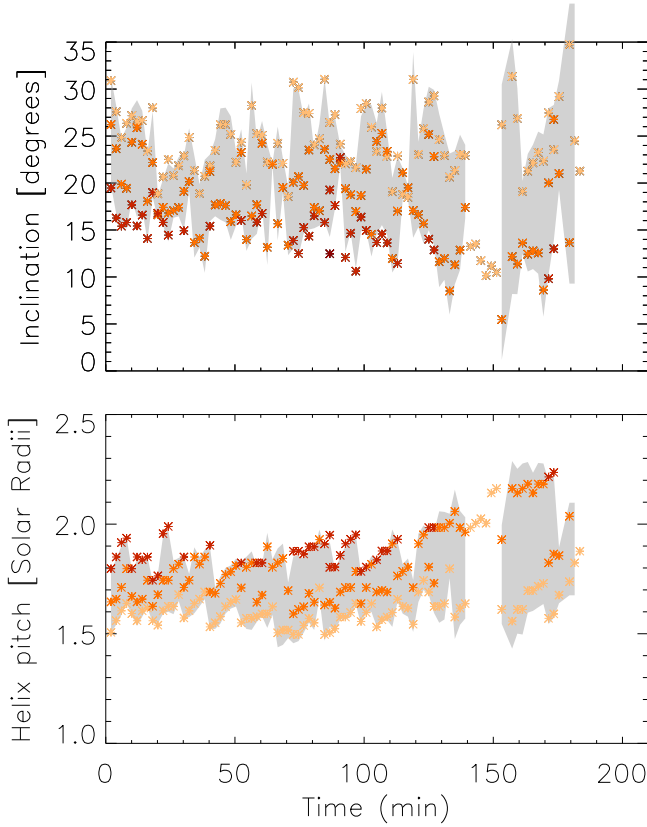
computational domain extends farther in terms of radial distance, we observe a slightly more scattered distribution.

Given that the goal of the simulation was a basic investigation of interchange reconnection and was in no way an “event study,” the similarity between the Metis observations and the MHD simulation is remarkable in several aspects. Notably, the persistent inclined features observed across different frames exhibit similar characteristics in terms of their evolution and apparent propagation from one frame to the next, despite the difference in size between the outflowing regions visible in the Metis images and that produced in the simulation. Additionally, we observe that the initial configuration of the MHD simulation resulted in an outflowing region with a substantially radial main axis, similar to what was observed by Metis, but this shifted laterally relative to the closed-field region, due to the twist injected by the driving. In this regard, we remark that the propagation direction of the polar crown prominence observed by EUV manifests a different overall inclination in comparison to the outflows visible in the Metis FOV.

## 6. Conclusions

In this paper, we present observations by Metis during its perihelion passage of a striking helical radial structure that extended from 1.5 to 3  $R_s$  and lasted for more than 3 hr. Thanks to the high resolution of the images and a long-duration data set (about 10 hr), with time sampling of 2 minutes and 1 s, we were able to measure the evolution of the helical structure and its fine-scale features. To the best of our knowledge, these



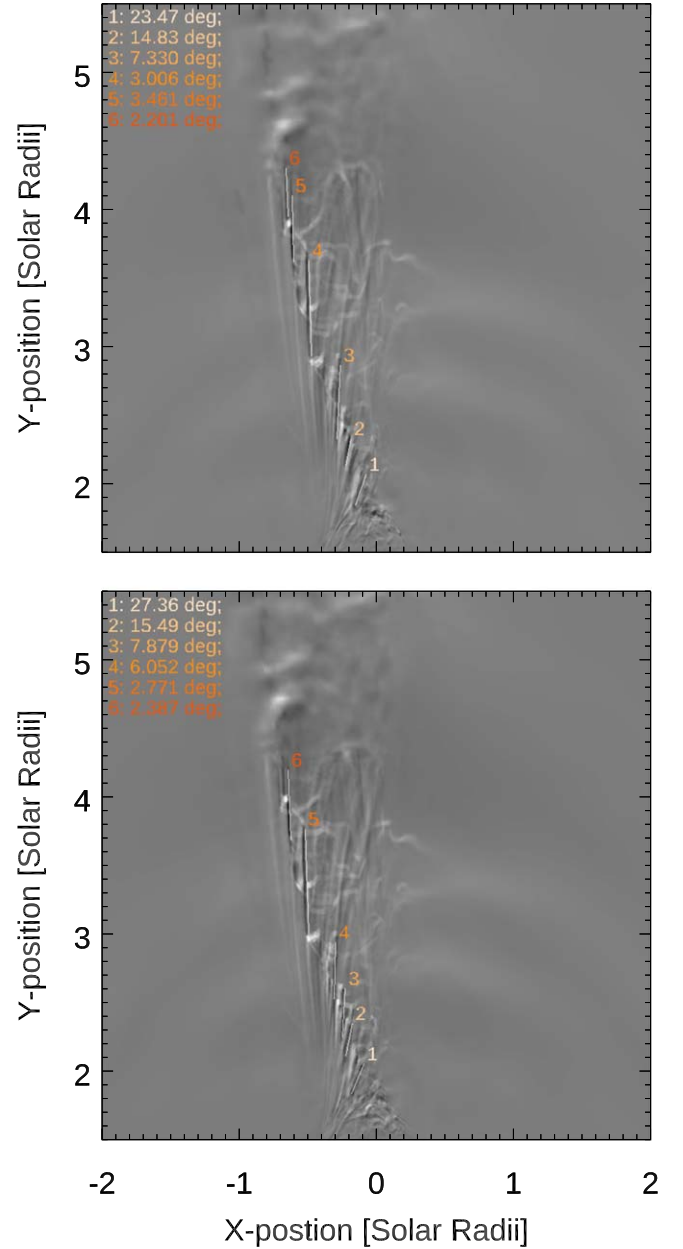


**Figure 11.** Temporal evolution of the inclination and helix pitch of the helicoidal structures, as observed by Metis from 10:15:01 UT ( $t=0$ ) to 13:36:56 UT on 2022 October 12. The gray regions represent the values corresponding to the average of the quantity  $\pm$  the standard deviation.

observations are unique, in that they appear to show directly the long-duration outflow of Alfvénic solar wind into the heliosphere. The observed features forming the helical structure are likely the result of bursty dense outflows and closed-field twist launched by interchange reconnection occurring in the wake of a jet-like CME, as indicated by the corresponding EUI images. An important finding is the long duration of the helical outflows, which persist until well after the initial eruption shown by EUI. We have shown that this can be explained naturally by the relaxation process expected within the framework of the breakout jet model in the topology of a large pseudostreamer (P. F. Wyper et al. 2018).

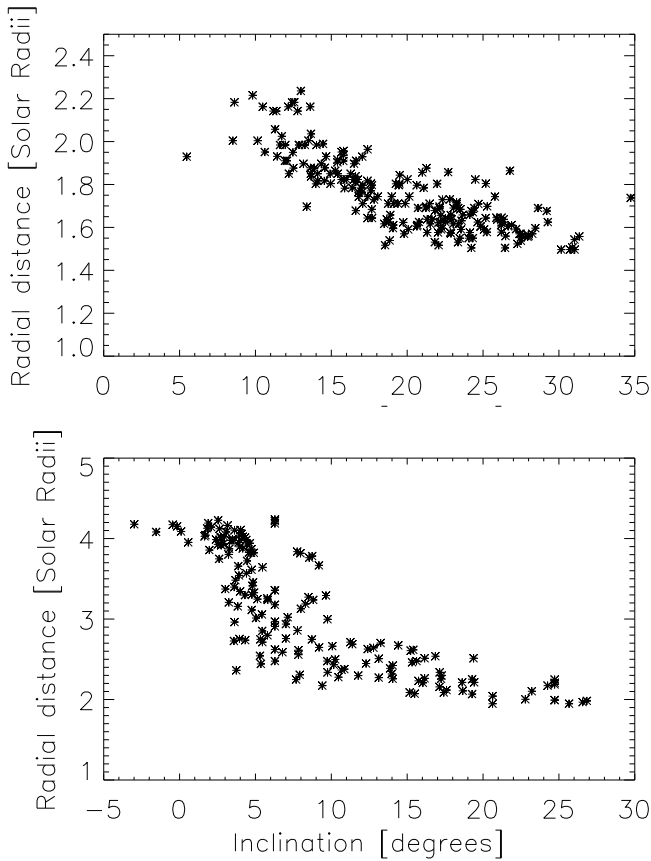
The spatial and temporal variation of the helical pitch and the inclination of the visible features allow us to interpret the helical structure as an open magnetic flux tube, characterized by an untwisting configuration of the magnetic field and corresponding to the stalk of the pseudostreamer. The fact that features located at higher polar angles and higher radial distances exhibit lower inclinations suggests that the field lines along which these features flow are becoming more radially aligned with distance and/or exhibit rotational motion across the plane of the sky and then into the plane. Additionally, the presence of bright and dark features northward of the easternmost helical features suggests a large portion of the pseudostreamer is involved with a fan-like outflow following its open fan-like spine.

The comparison between the Metis observations and the MHD simulations of P. F. Wyper et al. (2022) reveals striking similarities, particularly in the geometric parameters and



**Figure 12.** Two synthetic white-light images obtained by the P. F. Wyper et al. (2022) simulation. A measurement of the inclination of the inclined features and their height, similar to the one performed for the observations, was conducted for the simulations. In the top left of each panel, the angle between each structure and the radial direction is indicated.

evolution of these helical features. The magnetic field extrapolation, which detected a pseudostreamer at the same location as the helical structure and in agreement with the magnetic configuration of the simulation, reinforces the validity of this comparison. These similarities suggest that the observed features and their persistence for several hours are the result of long-duration interchange reconnection, consistent with the simulation. Furthermore, the characteristic black and white stripes in the synthetic white-light running-difference images derived from the simulation demonstrate a new and straightforward way to identify the signature of bursty interchange reconnection, which, due to the high time cadence of Metis, appear to have been observationally verified we believe for the first time.



**Figure 13.** Radial distance of the features vs. inclination as determined from the Metis observations (top) and the MHD simulation (bottom).

Furthermore, considering that the P. F. Wyper et al. (2022) simulation provides a potential explanation for the origins of the waves and fast outflows that evolve to produce the observed strong drops or reversals of the radial magnetic field component in the solar wind, accompanied by spikes in radial velocity (magnetic switchbacks; S. D. Bale et al. 2019; J. C. Kasper et al. 2019), we expect that events similar to the one reported in this work are critical for understanding the formation, evolution, and eventual dissipation of localized magnetic structures in the solar wind. One possibility is that the reconnection processes underlying our unique observations could serve as sources of perturbations that, as they propagate, become amplified or give rise to instabilities within the solar wind. Our observations, therefore, might be revealing the origins of the switchbacks/microstreams observed by the PSP mission, but many more studies are needed to clarify these connections.

In conclusion, the observations made by Metis during its perihelion passage provide unique and important insights into the basic dynamics shaping the outer corona and inner heliosphere. The consistency between these observations and the MHD simulations demonstrates that interchange reconnection is the crucial process for the formation of helical structures that are signatures of torsional Alfvén waves and solar wind outflows. We emphasize that our observations and simulation are generic, in that they capture the fundamental process of

interchange reconnection associated with the release of magnetic stress in the closed-field region of an embedded bipole. The only significant physical difference between our jet CME event and the ubiquitous jets and microjets that have been proposed as the source of the Alfvénic wind is spatial scale. Our results imply that long-lived Alfvénic outflows should be a universal feature of the smaller jets, as well. On the other hand, spatial scale is likely to be highly significant in determining whether the outflows contribute to the wind or simply fall back down to the chromosphere. The outflows that we observe start high up and appear to escape the Sun, but this may not be the case for the outflows from very-low-lying jets in the network. We conclude, therefore, that a definitive understanding of the generation of the Alfvénic solar wind requires many further studies. This conclusion underscores the importance of future high-resolution, long-duration coronal observations for advancing our knowledge of solar wind origin and evolution.

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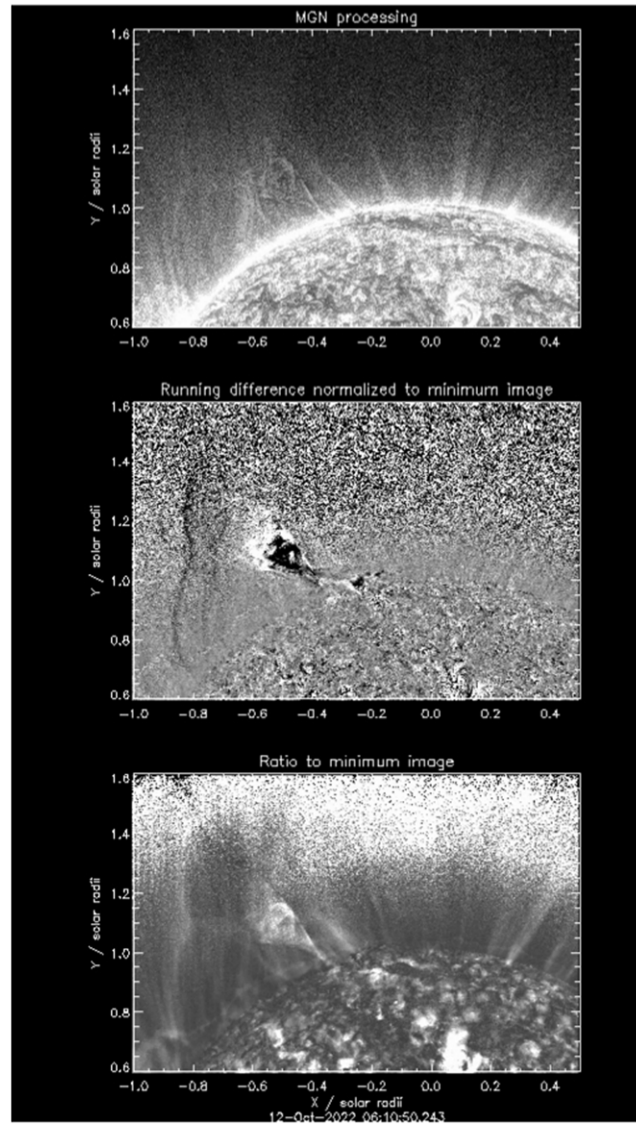
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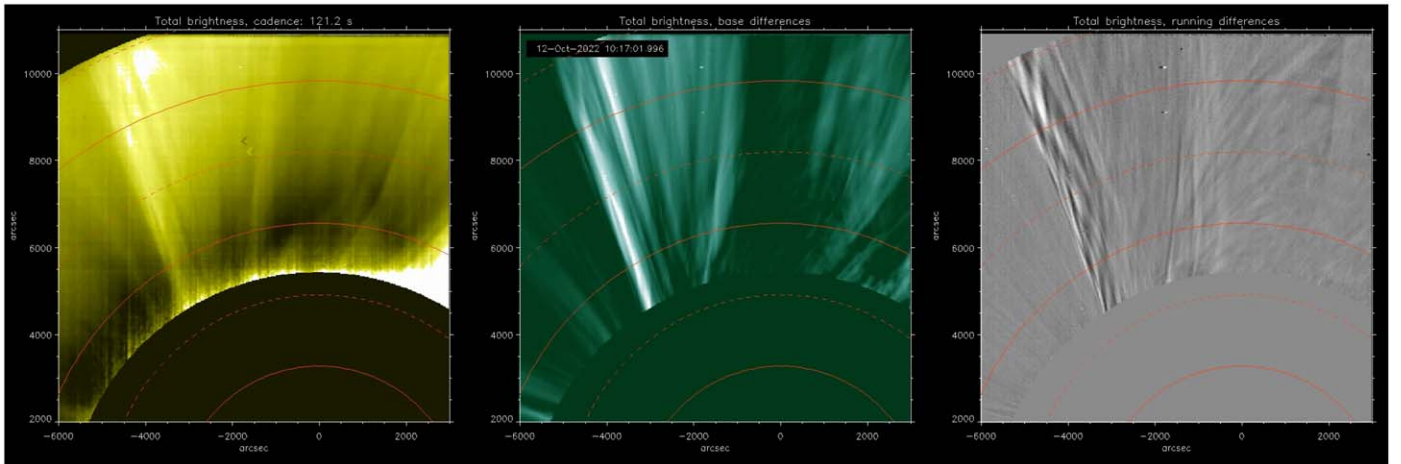
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### Appendix

Two supplemental figures, Figures 14 and 15, and their associated animations are provided, which show the EUI map and the Metis data set evolution.



**Figure 14.** A still image from an animation covering the EUV observations from October 12 at 04:00 UT to October 13 at 00:30 UT. Top: the EUV images are processed with Multiscale Gaussian Normalization (H. Morgan & M. Druckmüller 2014). Middle: the running-difference-normalized-to-minimum image. Bottom: the ratio-to-minimum image. In the movie, a linear feature becomes visible ahead of the erupting prominence at around 5:20 UT. The real-time duration of the movie is 24 s. (An animation of this figure is available in the online article.)



**Figure 15.** A still image from an animation showing the persistence of the helical features in the first 3.5 hr of observation of the Metis data set, taken from from 10:15 UT to 20:16. The left panel shows the total-brightness evolution. The middle panel is the total brightness with base differences. The total-brightness sequence with running differences is shown in the right panel. The animation has a time cadence of 2 minutes and 1 s and its real-time duration is 15 s. (An animation of this figure is available in the online article.)



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